

Discovery and follow-up studies of the extended, off-plane, VHE gamma-ray source HESS J1507-622

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ABSTRACT

Context. The detection of gamma-rays in the very-high-energy (VHE) range (100 GeV–100 TeV) offers the possibility of studying the parent population of ultrarelativistic particles found in astrophysical sources, so it is useful for understanding the underlying astrophysical processes in nonthermal sources.

Aims. The discovery of the VHE gamma-ray source HESS J1507-622 is reported and possibilities regarding its nature are investigated.

Methods. The H.E.S.S. array of imaging atmospheric Cherenkov telescopes (IACTs) has a high sensitivity compared with previous instruments (~1% of the Crab flux in 25 h observation time for a 5σ point-source detection) and has a large field of view (~5° in diameter). HESS J1507-622 was discovered within the ongoing H.E.S.S. survey of the inner Galaxy, and the source was also studied by means of dedicated multiwavelength observations.

Results. A Galactic gamma-ray source, HESS J1507-622, located ~3.5° from the Galactic plane was detected with a statistical significance $>9\sigma$. Its energy spectrum is well fitted by a power law with spectral index $\Gamma = 2.24 \pm 0.16_{\text{stat}} \pm 0.20_{\text{sys}}$ and a flux above 1 TeV of $(1.5 \pm 0.4_{\text{stat}} \pm 0.3_{\text{sys}}) \times 10^{-12} \text{ cm}^{-2} \text{ s}^{-1}$. Possible interpretations (considering both hadronic and leptonic models) of the VHE gamma-ray emission are discussed in the absence of an obvious counterpart.

Key words. gamma rays: general – X-rays: general

1. Introduction

In the extension of the successful H.E.S.S. survey of the Galactic plane (Aharonian et al. 2005, 2006a), performed in 2006/2007 with the High Energy Stereoscopic System (H.E.S.S.), a number

of new sources have been discovered and many of them are still unidentified (Aharonian et al. 2008). HESS J1507-622 is among the brightest (~8% of Crab Flux) newly discovered sources, and it lacks plausible counterparts, as is the case for HESS J1427-608, HESS J1708-410 (Aharonian et al. 2008), and HESS J1616-508 (Aharonian et al. 2006a). While all unidentified VHE sources that have been discovered in the H.E.S.S.

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Galactic plane scan are so far located within ± 1 degree from the Galactic equator, HESS J1507-622 is unique in this respect since it lies $\sim 3.5^\circ$ from the Galactic plane. Considering the comparably low hydrogen column density n_H at 3.5° off the plane, hence the lower Galactic absorption in X-rays and the reduced background emission, one would expect to detect a bright counterpart despite the anticipated lower spatial source density of Galactic counterparts.

Most Galactic VHE emitters are connected to young stellar populations, usually concentrated near the Galactic disk. Therefore, it is quite surprising to find an unidentified VHE gamma-ray source with a 3.5° offset from the Galactic plane, unless the source is nearby. HESS J1507-622 may offer the intriguing possibility of constraining the environment of an object that emits very-high energy radiation but lacks obvious counterparts at lower energies. Our line of sight towards the source direction intersects the Galactic disk with a scale height of 50 pc up to a distance of, at most, 1 kpc extending farther into the Galactic halo.

2. Observations

2.1. H.E.S.S. observations and data analysis

Within the 2006/2007 extension of the H.E.S.S. survey of the Galactic plane there were indications of a source at the position of HESS J1507-622; follow-up dedicated observations were performed for 9.7 h (21 runs of ~ 28 min each) in 2007 and in 2008. The data are calibrated as detailed in Aharonian et al. (2004). A thorough discussion of the H.E.S.S. standard analysis and performance of the instrument can be found in Aharonian et al. (2006c).

The peak significance, calculated following the method of Li & Ma (1983), is 9.3σ for the 9.7 h of dedicated observations (using a 0.22° oversampling radius, which is the standard radius used in source searches in the H.E.S.S. Galactic plane survey). Figure 1 shows the uncorrelated excess count map (smoothed with Gaussian of 0.07°) for the dedicated observations, using *hard cuts* (Aharonian et al. 2006c) and the *ring background method* (Berge et al. 2007). A two-dimensional Gaussian fit yields the best position at RA = $226.72^\circ \pm 0.05^\circ$, Dec = $-62.35^\circ \pm 0.03^\circ$, and the source is slightly extended with intrinsic size (not including the PSF) of $0.15^\circ \pm 0.02^\circ$ radius. The peak significance for this smaller (0.12° , chosen to reveal the source morphology at a reasonable significance without over-smoothing) correlation radius is 7.0σ .

The energy spectrum is reconstructed (0.22° extraction radius, which is larger than the source, since the morphology of the source is not well known) with the method presented in Aharonian et al. (2006c) with background subtracted using the *reflected-region method* (Berge et al. 2007): the results are shown in Fig. 2. Using *standard cuts* (Aharonian et al. 2006c), hence a lower energy threshold (~ 500 GeV), the observed spectrum can be well fitted with a power-law $dN/dE = k(E/1 \text{ TeV})^{-\Gamma}$ with photon index $\Gamma = 2.24 \pm 0.16_{\text{stat}} \pm 0.20_{\text{sys}}$ and a flux normalization $k = (1.8 \pm 0.4) \times 10^{-12} \text{ TeV}^{-1} \text{ cm}^{-2} \text{ s}^{-1}$, making the integral flux (above 1 TeV) $(1.5 \pm 0.4_{\text{stat}} \pm 0.3_{\text{sys}}) \times 10^{-12} \text{ cm}^{-2} \text{ s}^{-1}$. Using *hard cuts* (energy threshold ~ 1 TeV) and hence a better gamma-hadron separation (Aharonian et al. 2006c), the observed spectrum is well fitted by a power-law with photon index $\Gamma = 2.49 \pm 0.18_{\text{stat}} \pm 0.20_{\text{sys}}$ and a flux normalization $k = (3.1 \pm 0.8) \times 10^{-12} \text{ TeV}^{-1} \text{ cm}^{-2} \text{ s}^{-1}$, yielding an integral flux (above 1 TeV) of $(2.1 \pm 0.6_{\text{stat}} \pm 0.4_{\text{sys}}) \times 10^{-12} \text{ cm}^{-2} \text{ s}^{-1}$. The data points are compatible and the difference in flux arises

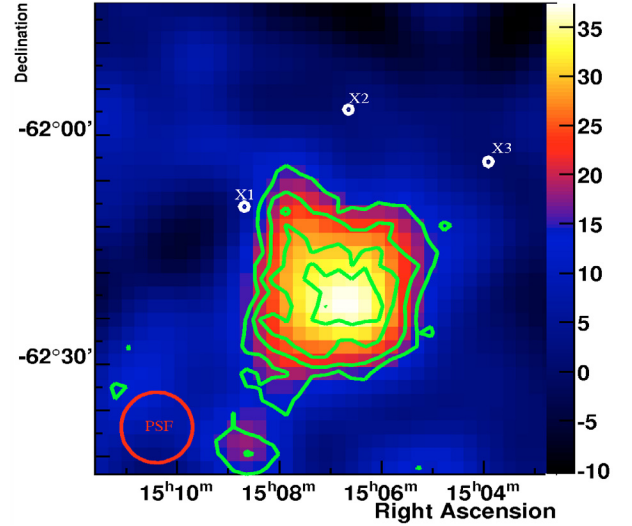


Fig. 1. Excess count map (smoothed with a Gaussian of 0.07° radius) of HESS J1507-622. The green contours show 3σ , 4σ , 5σ , 6σ contour levels for an integration radius of 0.12° . The white circles represent the position of three faint RASS (Voges et al. 2000) sources: X1 indicates 1RXS J150841.2-621006, X2 indicates 1RXS J150639.1-615704 and X3 indicates 1RXS J150354.7-620408. The 68% containment radius for the point spread function (PSF) of the H.E.S.S. instrument for these observations is shown in magenta.

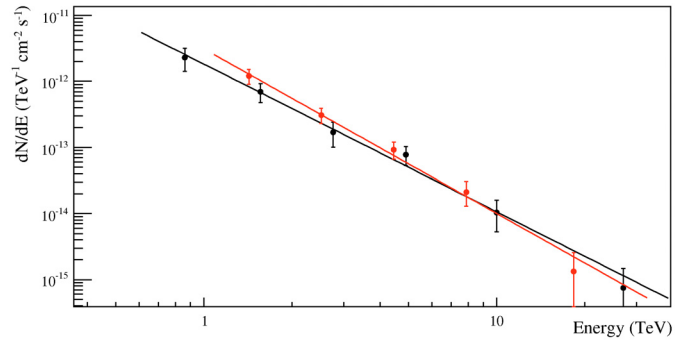


Fig. 2. The VHE spectrum observed from HESS J1507-622. The black line represents the best χ^2 fit of a power law to observed data using the *standard cuts*, whereas the red line shows the spectral result using *hard cuts* as described in Aharonian et al. (2006c).

from slope difference and from extrapolation to 1 TeV. No significant evidence of curvature in the spectrum has been found. An independent analysis and calibration based on a fit of camera images to a shower model (de Naurois 2006) gives compatible results.

2.2. X-ray observations and data analysis

The source was observed with XMM-Newton on Jan. 27, 2009 for 28 ks and with Chandra on Jun. 6, 2009 for 20 ks.

Unfortunately, the occurrence of a huge soft proton flare dramatically affected the 28 ks observation with XMM-Newton, leading to a good time interval (GTI) of only 0.8 ks for the PN detector, 8.0 ks for MOS1 and 9.2 ks for MOS2 detector. The soft proton flare is identified by extracting the high-energy (> 10 keV) light curve of the whole observation and is confirmed by the radiation monitor onboard XMM-Newton. However, background proton events still affect the remaining GTI, as they lead to wrong vignetting correction, in particular in the outer

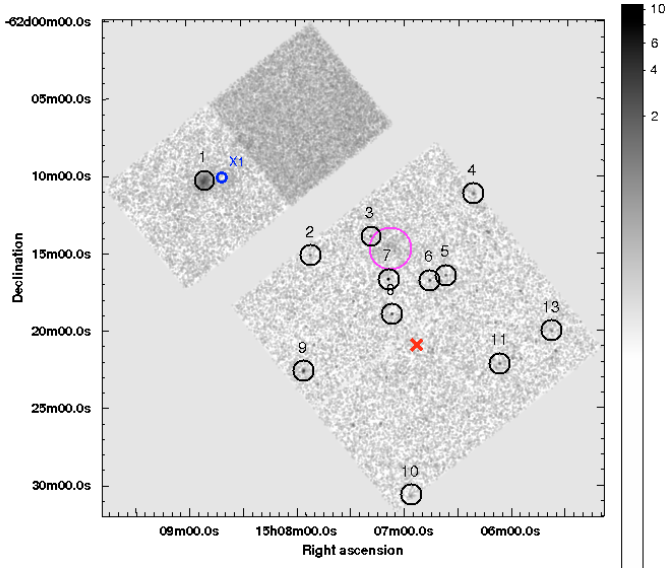


Fig. 3. Smoothed and background-subtracted count map of Chandra observation of HESS J1507-622. Black circles indicate the 12 sources detected by `celldetect`. The blue circle indicates 1RXS J150841.2-621006, centered on its nominal position and with radius corresponding to its positional uncertainties (Voges et al. 2000). The faint extended emission described in the text is shown in magenta. The red cross indicates the centroid position of HESS J1507-622.

regions of the MOS detectors. The data were reprocessed and the GTI was analyzed with SAS 9.0, discovering one point-like source (source 7 in Fig. 3) with a flux of $(1.4 \pm 0.4) \times 10^{-13}$ erg cm $^{-2}$ s $^{-1}$ in the energy range 2–10 keV: a compatible flux value has been found within Chandra observations and the point-like source is identified as a star (Tyc 9024.1705.1).

Figure 3 shows the smoothed and background subtracted count map between 0.3 and 8 keV as seen by Chandra for 20 ks. Chandra data were analyzed using CIAO 4.1. The background was estimated using the standard blank sky field files, since HESS J1507-622 covers the whole ACIS-I field of view. Eleven point-like sources (mainly identified as stars) and one extended source have been discovered using the source detection algorithm `celldetect`.

One of the most interesting discoveries by this Chandra observation concerns 1RXS J150841.2-621006 (source 1 in Fig. 3). This source is extended (35–40 arcsec) and spatially coincident with the radio source MGPS J150850-621025 (Green et al. 1999). The source is detected with an energy flux of $(7.0 \pm 0.7) \times 10^{-13}$ erg cm $^{-2}$ s $^{-1}$ in the range 2–10 keV, compatible with the flux measured by ROSAT. Even if the position of 1RXS J150841.2-621006 (Voges et al. 2000) does not coincide with the position of the extended Chandra source, the two positions are compatible when considering the well known systematics in the positional accuracy of ROSAT (Harris et al. 1998). The nature of the extended source 1RXS J150841.2-621006 is still unclear, since its spectrum is well fitted ($\chi^2/\text{d.o.f.} = 30.4/36$) by a simple power law (spectral index $= 2.0 \pm 0.3$; $n_H = (8.2 \pm 2.7) \times 10^{21}$ cm $^{-2}$), but the fit is equally good ($\chi^2/\text{d.o.f.} = 34.7/36$) for a thermal plasma scenario ($kT = 5_{-2}^{+5}$ keV; $n_H = (7.0 \pm 2.7) \times 10^{21}$ cm $^{-2}$). However, given its offset from the VHE source emission peak, outside the 3σ significance contour, 1RXS J150841.2-621006 is not a plausible counterpart for our source.

The other important aspect of the Chandra analysis is the existence of a faint extended source that escaped the source

detection algorithm `celldetect` but is detected (116 counts over a level of background of 32 counts) by `vtpdetect`, which performs better for faint diffuse structures. This source seems extended (20–25 arcsec), and it is slightly above the background level. Its flux is $(1.1_{-0.5}^{+0.3}) \times 10^{-13}$ erg cm $^{-2}$ s $^{-1}$ in the energy band between 2 and 10 keV, resulting from a power-law fit to the spectrum extracted, using `specextract`, from a region of 50 arcsec radius, centered on RA $_{J2000} = 15^{\text{h}}07^{\text{m}}06.8^{\text{s}}$, dec $_{J2000} = -62^{\circ}14'45.0''$. This faint source is confirmed by means of `wavdetect` with a significance of 6.9σ . It has a much smaller angular size than the VHE source, so an association is not obvious with current data, although the possibility remains that it is a bright part of a larger, weaker source. There is unfortunately no strong conclusion that can be made, but it could be linked to an old pulsar wind nebula (PWN) still visible at VHE. In fact the VHE PWNe sizes generally increase with pulsar age while the X-ray PWNe sizes show the opposite trend; moreover, for pulsars older than $\sim 10^3$ years the VHE PWNe are typically 100–1000 times larger than the sizes of the X-ray PWNe (while the difference is only a factor 2 for some younger PWNe, like the Crab Nebula), as shown by Kargaltsev & Pavlov (2010, e.g. HESS J1718-385). However, in the case of HESS J1507-622, there is no pulsar emission detected, as discussed in the following section.

3. Discussion

The unique feature that distinguishes HESS J1507-622 from any other unidentified VHE source is its angular offset ($\sim 3.5^\circ$) from the Galactic plane. Such an offset may mean that HESS J1507-622 is a nearby, local object or that it is truly distant in the halo. This second possibility may have important implications for any gamma-ray production scenario. At a location away from the Galactic disk, the density of target material is quite low (Lockman 1984), making a hadronic scenario for HESS J1507-622 less favorable. On the other hand, the cosmic microwave background (CMB), as target photon field for gamma-ray production, is available everywhere pointing towards a leptonic gamma-ray emission mechanism. However, if confirmed by deeper observations, the lack of emission at other wavebands could imply a hadronic origin for HESS J1507-622. Both cases are discussed in this section.

3.1. Searching for counterparts

The search for counterparts was made following the procedure outlined in Aharonian et al. (2008). At radio and infrared wavelengths, HESS J1507-622 is offset too far from the Galactic plane to be covered by the Southern Galactic Plane Survey (Haverkorn et al. 2006) or by Spitzer GLIMPSE (Benjamin 2005). This region of the sky had been covered by the Midcourse Space Experiment (MSX) in all its wave bands (8.28 μm , 12.13 μm , 14.65 μm , 21.3 μm , Simon et al. 2006) and by the MOLONGLO Galactic plane survey (Green et al. 1999), yielding no evidence of any plausible counterpart. HESS J1507-622 is located on a radio emission filament shown in Duncan et al. (1995) at 2.4 GHz, which was tentatively considered a part of a very large ($\sim 15^\circ$ in diameter) and nearby candidate SNR (visible in Fig. 4). In the complete CO survey (Dame et al. 2001), the H.E.S.S. source lies near the edge of a large ($\sim 5^\circ \times \sim 2^\circ$) nearby CO molecular cloud (visible in Fig. 5), and the peak velocity of this cloud, around -5 km s $^{-1}$, would most likely place it quite near at a distance of ~ 400 pc. The substantial difference

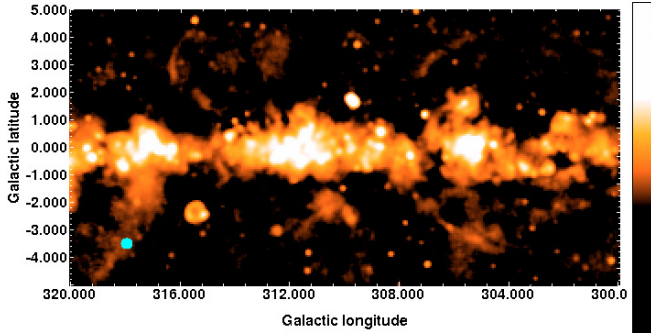


Fig. 4. 2.4 GHz radio map from Duncan et al. (1995). HESS J1507-622 is marked by a cyan circle centered at the nominal position of the VHE gamma-ray source and with 0.15° radius.

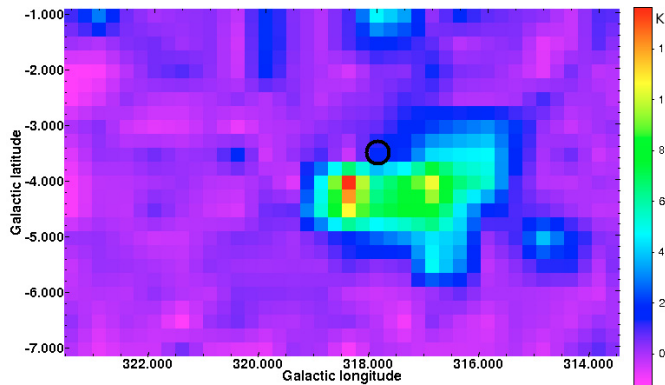


Fig. 5. CO map of the region around HESS J1507-622 integrated in the velocity parameter in the range -10 to 0 km s^{-1} ; the CO peak is at ~ -5 km s^{-1} , which corresponds to a distance of ~ 400 pc (or, more unlikely 12 kpc). The black circle represents HESS J1507-622: it is centered at the nominal position of the VHE gamma-ray source and has a 0.15° radius.

in extension and, in the case of the CO molecular cloud, the offset of $\sim 1^\circ$ from the H.E.S.S. source centroid, suggest no obvious scenario for an association.

The absence of obvious counterparts in X-rays in combination with the observed VHE flux can be used to compare the multiwavelength properties of HESS J1507-622 with other unidentified sources. The ratio of $F_\gamma(1-10 \text{ TeV})/F_X(2-10 \text{ keV})$ has been suggested as a viable parameter for such a comparison (Yamazaki et al. 2006). If the flux of the Chandra faint extended source (3) is adopted as an upper limit for any low-energy counterpart and using the H.E.S.S. flux found with hard cuts for HESS J1507-622, this ratio is estimated to be >100 . This would place the source above the ratio of F_γ/F_X for the first unidentified source of the H.E.S.S. Galactic plane scan HESS J1303-631, which was estimated to be >1.6 . Moreover, this ratio for HESS J1507-622 would be greater than the value obtained for the “darkest” VHE source so far of >55 for HESS J1616-508 (Matsumoto et al. 2007). The caveat of determining the ratio F_γ/F_X in this way is that the flux of the Chandra blob (extraction radius of $\sim 50''$) is compared to the VHE flux of a more extended object ($\theta \sim 0.15^\circ$).

If the previously discussed association of the faint extended X-ray source and HESS J1507-622 cannot be maintained, e.g. as the result of deeper or complementary multifrequency observations, we could also be faced with the case of not having any counterpart at all. In this case, HESS J1507-622, according to its position and to its brightness in VHE gamma-rays, could be a candidate for a “dark accelerator”, i.e. of a purely hadronic

source, visible mainly at TeV energies. On the other hand, it will also be argued that a mature PWN can appear to look like a “dark accelerator”; in this case, the emission would be from inverse Compton scattering, coming from a TeV calorimeter around the aging pulsar (see e.g. de Jager 2008).

3.2. Leptonic scenario: PWN

In a leptonic scenario it is expected that the same population of high-energy electrons that generate the gamma-ray emission should also produce nonthermal X-rays due to synchrotron radiation caused by the ambient magnetic field. The non-detection of X-ray emission from the location of HESS J1507-622 would therefore place limits on the strength of any magnetic field at the source. If we use the X-ray brightness of 1RXS J150841.2-621006 to place an upper limit on any synchrotron component connected to the detected VHE emission and if we assume that the gamma-ray emission is produced by electrons following a power-law distribution with a spectral index of -2.2 that up-scatter cosmic microwave background (CMB) photons, then the ambient magnetic field cannot exceed a strength of $0.5 \mu\text{G}$. The caveat mentioned in the previous section applies here since we have determined the magnetic field by comparing the flux of a ROSAT point source with the VHE flux of an extended object. For higher upper limits for any diffuse X-ray emission, the limit on the magnetic field is less constraining.

PWNe are usually considered natural candidates for leptonic Galactic gamma-ray emitters (e.g. Gaensler & Slane 2006). Although they are, in many cases, detected as nonthermal X-ray sources, in the following we outline how an evolved PWN can indeed lead to a fairly bright gamma-ray source without any counterpart (de Jager 2008). The key issue for such a model is that the low energy synchrotron emission, where τ_E is the synchrotron emitting lifetime of TeV gamma-ray emitting leptons with energy E , depends on an internal property of the PWN, namely the magnetic field (e.g. Kennel & Coroniti) which may vary as a function of time, following $\tau_E \propto t^{2\alpha}$ if $B(t) \propto t^{-\alpha}$, where α is the power-law index of the decay of the average nebular field strength, whereas the VHE emission depends on the CMB radiation field, which is constant on timescales relevant for PWN evolution. Magnetohydrodynamical (MHD) simulations of composite SNRs with different spindown power at birth and spindown timescale (de Jager et al. 2009) find that $\alpha = 1.3$ until the passage of the reverse shock (at time $t = T_R$), when $B(t)$ flattens out, or even increases slightly. As the simulation time only extended up to 10 kyr (with $B(t) \sim B(T_R)$), it is expected that field decay would continue for $t \gg T_R$ since expansion continues even after the passage of the reverse shock.

From hydrodynamic simulations, an expression for the time of the reverse shock passage (i.e. the return time of the reverse shock to the origin) has been given (Ferreira & de Jager 2008) as

$$T_R = 10 \text{ kyr} \left(\frac{\rho_{\text{ISM}}}{10^{-24} \text{ g/cm}^3} \right)^{-1/3} \left(\frac{M_{\text{ej}}}{10 M_\odot} \right)^{3/4} \left(\frac{E_{\text{ej}}}{10^{51} \text{ erg}} \right)^{-2/3}$$

where ρ_{ISM} is the density of the interstellar medium, E_{ej} the SNR blast wave energy and M_{ej} the ejecta mass during the SNR explosion. The stellar wind of a high-mass star can blow a cavity around the progenitor star (Chevalier & Liang 1989) with relatively low ISM density, so that $T_R \gg 10$ kyr. In such a case it is expected that $B(t)$ can decay as $t^{-\alpha}$, until the field is low enough for the X-ray flux to drop below the typical sensitivity levels, and that $t \gg \tau_0$ (where τ_0 is the spindown timescale), while $t < T_R$, so that the time integral of the total measured

Table 1. Summary of the twelve surces detected by Chandra using `celldetect`.

src	CXOU name	RA (J2000)	Dec(J2000)	Countrate	Possible identification	2MASS sources
1	CXOU J150850.6-621018	15:08:50.637	−62:10:18.24	$(4.6 \pm 0.2) \times 10^{-2}$	1RXS J150841.2-621006	
2	CXOU J150752.0-621509	15:07:52.091	−62:15:09.96	$(1.0 \pm 0.3) \times 10^{-3}$	3UC 056-316462 GSC2.3 S7QP004587 GSC2.3 S7QP084980	15075219-6215107 15075188-6215057 15075290-6215077
3	CXOU J150718.5-621357	15:07:18.516	−62:13:57.22	$(1.4 \pm 0.3) \times 10^{-3}$	3UC 056-316221 GSC2.3 S7QP005118	15071853-6213571
4	CXOU J150621.7-621110	15:06:21.789	−62:11:10.76	$(1.7 \pm 0.4) \times 10^{-3}$	TYC 9024.1615.1 3UC 056-315763	15062140-6211110
5	CXOU J150636.9-621628	15:06:36.935	−62:16:28.47	$(1.0 \pm 0.3) \times 10^{-3}$	–	–
6	CXOU J150645.7-621648	15:06:45.788	−62:16:48.10	$(2.5 \pm 0.4) \times 10^{-3}$	GSC2.3 S7QP084624 GSC2.3 S7QP084624	15064587-6216466
7	CXOU J150708.8-621643	15:07:08.824	−62:16:43.99	$(7.6 \pm 0.8) \times 10^{-3}$	TYC 9024.1705.1 3UC 056-316134	15070879-6216441
8	CXOU J150706.7-621858	15:07:06.746	−62:18:58.75	$(2.5 \pm 0.4) \times 10^{-3}$	GSC2.3 S7QP060692	15070693-6218560
9	CXOU J150756.0-622238	15:07:56.018	−62:22:38.70	$(4.2 \pm 0.7) \times 10^{-3}$	GSC2.3 S7QP035739 GSC2.3 S7QP035813	15075554-6222336 15075683-6222411
10	CXOU J150656.1-623040	15:06:56.128	−62:30:40.72	$(1.7 \pm 0.4) \times 10^{-3}$	GSC2.3 S7QP031378 GSC2.3 S7QP031375 GSC2.3 S7QP031476 GSC2.3 S7QP001419 GSC2.3 S7QP031426 GSC2.3 S7QP031294	15065496-6230425 15065736-6230444 15065721-6230287 15065613-6230553 15065368-6230417 15065363-6230347 15065746-6230579
11	CXOU J150606.7-622210	15:06:06.799	−62:22:10.36	$(1.7 \pm 0.4) \times 10^{-3}$	–	15060774-6222092
12	CXOU J150538.1-622001	15:05:38.114	−62:20:01.35	$(1.3 \pm 0.3) \times 10^{-3}$	GSC2.3 S7QP003678	15053757-6220015

Notes. For source 1 only the firm identification 1RXS J150841.2-621006 (Voges et al. 2000) is mentioned. The Tycho-2 catalogue (Hog et al. 2000), the second-generation guide star catalog (Lasker et al. 2008), the UCAC3 Catalogue (Zacharias et al. 2010), and the 2MASS All-Sky Catalog of Point Sources (Cutri et al. 2003) have been inspected when searching for possible counterparts.

energy of nebular leptons is convergent; i.e., the initial kinetic energy at birth of the pulsar is measured. As a result, in a scenario where the magnetic field decays as a function of time, the synchrotron emission will also fade as the PWN evolves. The reduced synchrotron losses for high-energy electrons for such a scenario will then lead to increased lifetimes for these leptonic particles. For timescales shorter than the inverse-Compton lifetime of the electrons ($t_{\text{IC}} \approx 1.2 \times 10^6 (E_e/1 \text{ TeV})^{-1}$ years), this will result in an accumulation of VHE electrons which will also lead to an increased gamma-ray production due to up-scattering of CMB photons. Such accumulation of very-high energy electrons in a PWN has indeed been seen in the source HESS J1825-137 (Aharonian et al. 2006d). To summarize, during their evolution PWN may appear as gamma-ray sources with only very faint low-energy counterparts and this may represent a viable model for HESS J1507-622.

Within this scenario, the faint diffuse emission detected by Chandra assumes a particular importance. The morphology of HESS J1507-622 is similar to HESS J1702-420 (Aharonian et al. 2008), where the pulsar PSR J1702-4128 could be powering an extremely asymmetric PWN visible in VHE gamma-rays, but it should be emphasized that the PWN association of HESS J1702-420 is not established. In the ancient PWN scenario therefore the faint diffuse emission in X-ray could be due to a younger population of electrons and reflect the proper motion of the pulsar within the nebula. Also the $F_\gamma(1-10 \text{ TeV})/F_X(2-10 \text{ keV})$ is similar in these two VHE unidentified sources; in fact, Fujinaga et al. (2009) calculated $F_\gamma(1-10 \text{ TeV})/F_X(2-10 \text{ keV}) > 32$ for HESS J1702-420.

In the framework of a PWN interpretation, it is also interesting to compare HESS J1507-622 with other VHE PWNe and PWN candidates that are located at a similar angular offset from the Galactic plane. These objects are (sorted by decreasing offset) Crab ($b = -5.8^\circ$) (Weekes et al. 1989), Geminga ($b = 4.3^\circ$) (Abdo et al. 2007), Vela X ($b = -3.1^\circ$) (Aharonian et al. 2006b) and HESS J1356-645 ($b = -2.5^\circ$) (Renaud et al. 2008). All these objects related to prominent pulsars and to their associated multiwavelength counterparts. Two of them are nearby sources (Geminga with $d = 160 \text{ pc}$, Caraveo et al. 1996; and Vela with $d = 290 \text{ pc}$, Caraveo et al. 2001) and feature very extended VHE emission regions (Geminga $\sim 2.8^\circ$, as indicated in Abdo et al. 2007, and Vela X $\sim 0.5^\circ$ radius, as measured by H.E.S.S.), as can be expected for close, extended objects. Geminga is the only example here of an old pulsar: its age has been estimated to be $\sim 300\,000$ years (Bertsch et al. 1992). In comparison to all these sources, HESS J1507-622 is different since it lacks any obvious counterpart. If it is indeed an evolved PWN, then its rather compact appearance would disfavor a very small distance; by simple comparison with Geminga, we should conclude that the PWN scenario constrains HESS J1507-622 distance to a value $> 6 \text{ kpc}$.

3.3. Hadronic scenario: SNR and GRB remnant

For the hadronic scenario of gamma-ray production, the total energy in cosmic rays in the object can be derived from the gamma-ray luminosity and the density of target material. Since

HESS J1507-622 is located off the Galactic disk, the density of target material is lower than in the Galactic plane and can be obtained as a function of distance, following the density profile of the interstellar medium perpendicular to the Galactic plane. Using the best-fitting model (Lockman 1984) for the density profile off the Galactic plane at relevant galactocentric distances between 4 and 8 kpc and assuming a spectral index of -2 and a cut-off at 100 TeV for the cosmic rays, it is found that about 10^{50} ($\sim 10^{51}$) erg of cosmic rays at a distance of 2 (4) kpc are needed to explain the VHE emission observed in HESS J1507-622. The very fast increase in the energy in cosmic rays with distance (a factor of 10 for an increase in distance by a factor of 2) results from the fact that a larger distance places the source also at a larger spatial offset from the Galactic disk, hence in an environment with lower density of target material.

The main accelerators of hadronic cosmic rays are believed to be the shock waves of SNRs. If HESS J1507-622 is indeed connected to an SNR, due to its energetics, it would be located at a distance of $\lesssim 2$ kpc (density of 0.45 cm^{-3}) (Lockman 1984). A supernova remnant at a distance of 2 kpc with the apparent angular size of HESS J1507-622 would be in the pressure-driven Sedov phase with an age of less than 1000 years. It is interesting to note that such an SNR should feature similar observable properties as the remnant of SN 1006. However, SN 1006 is a prominent emitter of nonthermal X-ray emission, e.g. ASCA (Koyama et al. 1996) and ROSAT (Willingale et al. 1996), and the absence of such X-ray emission in HESS J1507-622 disfavors any SNR origin for this VHE source. In this scenario the faint X-ray extended emission could be part of an SNR shell.

A proposed more exotic class of hadronic VHE gamma-ray sources are remnants of gamma-ray bursts (Atoyan et al. 2006). If HESS J1507-622 is indeed the result of a GRB, it would be of the subclass of compact binary merger induced short GRBs since the location of HESS J1507-622 off the Galactic plane would indicate that it is connected to an old stellar population (Domainko & Ruffert 2005, 2008). Also the X-ray afterglows of short bursts indicate a low-density environment that would be consistent with a location away from galactic disks (see e.g. Nakar 2007). In the model of a GRB remnant, hadronic cosmic rays are produced in a point-like explosion and extended, center-filled VHE gamma-ray sources are then the result of energy dependent diffusion of the cosmic rays and subsequent hadronic gamma-ray production. In this picture the age of the remnant can be estimated from the source extension to $t/1 \text{ kyr} \lesssim 0.2 (d/1 \text{ kpc})^2$ assuming typical values of Galactic cosmic ray diffusion (Atoyan et al. 2006), which may be different outside the Galactic plane. The rate of merger-induced bursts in the Galaxy is found to be about one event every $(0.5-7) \times 10^4$ years (Kalogera et al. 2004; Guetta & Piran 2006). For an anticipated age of $\gtrsim 10^4$ years the remnant has to be at a distance of $\gtrsim 7$ kpc. At such a distance the total energy of cosmic rays would be $\sim 10^{52}$ erg, which is quite high for the energetics of typical short bursts (Nakar 2007). To summarize, a GRB remnant interpretation of HESS J1507-622 would require a highly energetic and relatively recent merger event ($\lesssim 10^4$ years) which may be possible, but is quite unlikely.

4. Conclusions

The discovery of HESS J1507-622, a rather bright ($\sim 8\%$ of the Crab flux) unidentified Galactic VHE gamma-ray source, is reported. The lack of obvious counterparts, together with its location 3.5° off the Galactic plane, play a crucial role in interpreting this source, which shows the highest value yet found for the ratio between X-ray and VHE emission. Hadronic and

leptonic scenarios for the VHE gamma-ray radiation are discussed, showing that a PWN scenario with an old pulsar could be a possible interpretation for HESS J1507-622, while a hadronic model (e.g., SNR-related hadronic emission and GRB remnant) appears unlikely unless the distance is fairly small. Upcoming deeper X-ray observations (XMM-Newton and Suzaku) will undoubtedly offer deeper insight in this VHE source.

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